

Ices in Star-Forming Regions: First Results from VLT-ISAAC

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Abstract. The first results from a VLT-ISAAC program on the infrared spectroscopy of deeply-embedded young stellar objects are presented. The advent of 8-m class telescopes allows high S/N spectra of low-luminosity sources to be obtained. In our first observing run, low- and medium-resolution spectra have been measured toward a dozen objects, mostly in the Vela and Chamaeleon molecular clouds. The spectra show strong absorption of H_2O and CO ice, as well as weak features at '3.47' and $4.62\ \mu m$. No significant solid CH_3OH feature at $3.54\ \mu m$ is found, indicating that the CH_3OH/H_2O ice abundance is lower than toward some massive protostars. Various evolutionary diagnostics are investigated for a set of sources in Vela.

1 Introduction

Interstellar matter provides the basic building blocks from which new solar systems like our own are made. The formation of stars and planetary systems begins with the collapse of a dense interstellar cloud core, a reservoir of dust and gas from which the protostar and circumstellar disk are assembled. In this cold and dense phase, molecules freeze-out onto the grains and form an icy mantle surrounding the silicate and carbonaceous cores. The ices can contain up to 40 % of the condensable material (i.e., C and O) (Whittet et al. 1998, d'Hendecourt et al. 1999). Much of this material is incorporated into the circumstellar disks and ultimately in icy solar system bodies such as comets (Ehrenfreund et al. 1997). A central quest in star formation and astrochemistry is to understand the evolution of these species from interstellar clouds to planetary bodies, and use them as diagnostic probes of the thermal history and physical processes (van Dishoeck & Blake 1998, Ehrenfreund & Charnley 2000, Langer et al. 2000).

In recent years, much progress has been made in our understanding of the chemical evolution during star formation through combined submillimeter and infrared observations. Submillimeter observations using telescopes such as the JCMT, CSO, IRAM 30m and SEST probe the gas-phase composition of the warm and dense envelopes around deeply-embedded protostars. Many different

species with abundances down to 10^{-11} with respect to H_2 can be probed through their pure rotational transitions. Because of the high spectral resolution of the heterodyne technique ($R = \lambda/\Delta\lambda > 10^6$), the line profiles are resolved and provide information on the location of the molecules (e.g., quiescent gas vs. outflow). Since the lines are in emission, a map of their distribution can be made.

Infrared spectroscopy provides complementary information: at these wavelengths the vibrational modes of both gas-phase and solid-state species can be observed, but only down to abundances of $\sim 10^{-7}$ with respect to H_2 at typical resolving powers $R \approx$ a few thousand. Symmetric molecules, such as H_3^+ , CH_4 , C_2H_2 and CO_2 have no dipole-allowed rotational transitions, and can therefore only be probed through their strong infrared vibrational transitions. Moreover, PAH emission features appear throughout the mid-infrared range, and the dominant interstellar molecule, H_2 , has its pure rotational transitions at these wavelengths. The cold gases and ices are usually observed in absorption toward an embedded YSO, where the hot dust in the immediate circumstellar environment provides the continuum. This technique samples only a pencil beam line of sight. The *Short Wavelength Spectrometer* (SWS) on the *Infrared Space Observatory* (ISO) has opened up the mid-infrared wavelength region and has obtained spectra of more than a dozen protostars without atmospheric interference (see van Dishoeck & Tielens 2001 for review). However, ISO was limited to the most luminous, massive YSOs ($L \approx 10^4 - 10^5 L_\odot$). A major goal of our VLT-ISAAC program is to extend this work to lower-mass objects ($L \leq 10^3 L_\odot$) representative of our proto-Sun.

Together, the submillimeter and infrared data have led to the following scenario for high-mass objects. In the cold pre-stellar cores and collapsing envelopes, gas-phase molecules freeze-out onto the grains and form an icy mantle. Here the abundances can be further modified by grain surface reactions and, perhaps, photoprocessing of ices. In particular, the hydrogenation and oxidation of accreted C, O, N and CO can lead to CH_4 , H_2O , NH_3 , H_2CO , CH_3OH and CO_2 , respectively (Tielens & Charnley 1997). Most of these species have been firmly identified in interstellar ices (e.g., Whittet et al. 1996, d'Hendecourt et al. 1996, Gibb et al. 2000). Once the protostar has formed, its luminosity can heat the surrounding grains to temperatures at which the ices evaporate back into the gas phase, resulting in enhanced gas/solid ratios (e.g., van Dishoeck et al. 1996, Dartois et al. 1998, Boonman et al. 2000). The sublimation temperatures range from ~ 20 K for pure CO ice to ~ 90 K for H_2O -rich ice under typical conditions. These freshly evaporated molecules can then drive a rich and complex chemistry in the gas (called the ‘hot core’ phase) until the normal ion-molecule chemistry takes over again after $\sim 10^5$ yr (Charnley et al. 1992). The different solid-state and gas-phase species therefore serve not only as physical diagnostics, but also as probes of the evolution of the region.

2 Our VLT-ISAAC program

In 1999, we proposed a large VLT-UT1 ISAAC program to probe the origin and evolution of ices in southern star-forming regions through a spectroscopic survey of 30 – 40 objects in the 2.7–5.1 μm L and M-band atmospheric windows. The aim was to obtain high-quality spectra ($S/N > 50$ on continuum) such that 3σ limits of species with abundances down to 2–4% of H_2O ice — the dominant ice component — can be obtained. The program focuses on low- and intermediate-mass southern YSOs, but covers a range of evolutionary stages from background stars to T-Tauri stars with disks. Also, several different environments (Vela, Chamaeleon, Ophiuchus, Corona Australis, ...) will be probed. The program was awarded 14 nights of VLT time. However, due to technical problems with the long-wavelength arm of ISAAC, the first observing session did not take place until January 2001.

The January 2001 observations totalled 5 nights and were carried out under mediocre conditions with high humidity. Nevertheless, good low-resolution (LR, $R = 600 - 800$) spectra were obtained for ~ 15 objects and medium-resolution (MR, $R = 3000 - 5000$) spectra for ~ 5 objects, covering \sim one-third of our project. With the new 1024×1024 Aladdin array, the low-resolution spectra can be obtained in a single spectral setting per atmospheric window, whereas the medium resolution spectra require 6/3 settings to cover the entire L/M band, respectively. Typical exposure times are ~ 30 minutes per atmospheric window in LR for a $L \approx 7$ mag object. Standard stars were observed immediately before or after the YSOs, within 0.05–0.1 airmass. The spectra were obtained using both chopping (by $15''$ along the slit) and nodding of the telescope. Daily arcs and flat-fields were provided by the observatory staff. In the M-band, the atmospheric CO lines were used for wavelength calibration. The data were reduced using IDL routines developed in-house by E. Dartois and W.F. Thi.

Because the ice features are weak and broad and are superposed on a strong continuum, many tests were carried out to check the reliability of the spectra. Specifically, spectra were taken for objects also observed with ISO or UKIRT, and features were checked on different nights at the VLT and between the LR and MR modes on the VLT. In general, the reproducibility of the features is excellent (see Fig. 1).

A particularly nice capability of ISAAC is the possibility to rotate the slit on the sky to obtain spectra of more than one object simultaneously. In several cases, more than one bright object at L or M-band was discovered in the acquisition image within $\sim 15''$. The slit was then rotated to include the additional object. An example is shown in Fig. 2 for the high-mass source GL 961, where the E and W components are separated by $\sim 5.5''$.

Laboratory data such as those obtained in the Sackler Laboratory for Astrophysics in Leiden and at the Astrochimie Expérimentale Laboratory in Paris play an essential role in the analysis of the infrared spectra. First, they lead to definite identification of the molecules and to quantitative estimates of their abundances. Second, they provide an indication of the ice environment and ice components. For example, a CO molecule embedded in an H_2O -rich matrix ('polar ice') has a

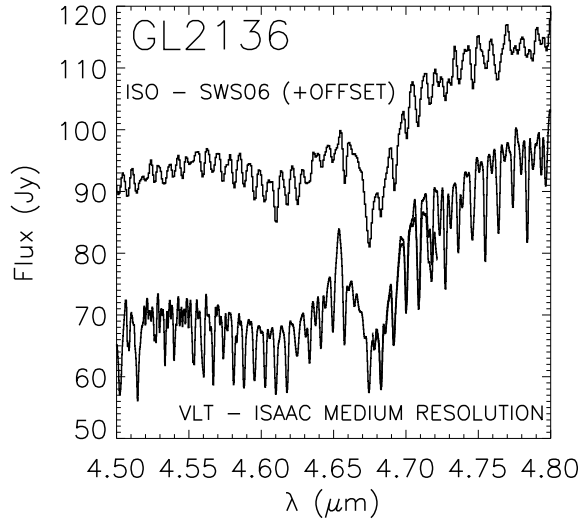


Fig. 1. VLT-ISAAC MR M-band $R \approx 2000$ spectrum of GL 2136 compared with the ISO-SWS spectrum ($R \approx 1500$). The sharp features are due to gas-phase CO, and become stronger in the higher-resolution VLT spectrum. Broad absorption features at $4.62 \mu\text{m}$ due to solid OCN^- and at $4.67 \mu\text{m}$ due to solid CO are seen as well. The H I 7–5 emission line at $4.6538 \mu\text{m}$ is stronger in the VLT spectrum due to the smaller slit.

different spectral shape compared with that in a CO-rich mantle (‘apolar ice’). Such analyses have shown that the ice mantles in the protostellar environment are not homogeneous, but consist of several components. These phases may reflect differential accretion of atomic H-rich versus H-poor gas and/or different degrees of outgassing of the more volatile species (Schutte 1999). Third, the ice spectra and the gas/solid ratios can provide clear evidence for heating of the ices in the more evolved objects (Ehrenfreund et al. 1998, Boogert et al. 2000, van Dishoeck et al. 1996).

3 Initial results

The prime targets in the January 2001 run were a set of low- and intermediate mass YSOs in the Vela and Chamaeleon clouds, together with a few well-known southern high-mass protostars. In the first case, the power of the VLT+ISAAC is used to observe weaker and lower luminosity sources than possible previously, whereas for the high-mass sources the aim was to obtain higher S/N and higher spectral resolution data than provided by ISO or previous ground-based observations.

VLT - ISAAC LW

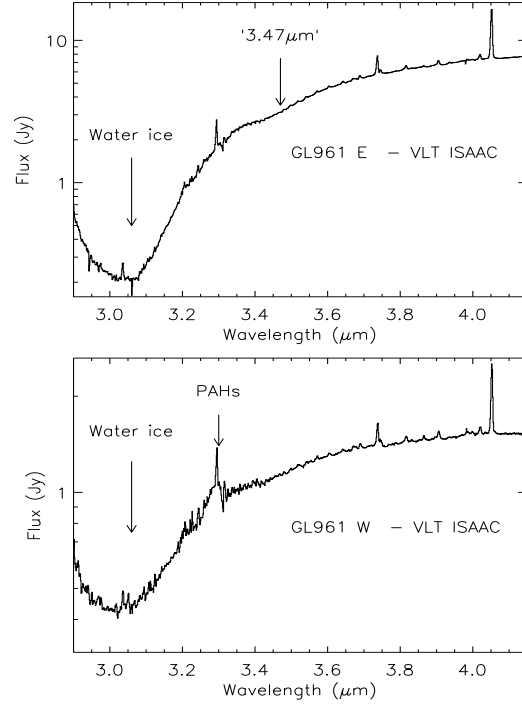
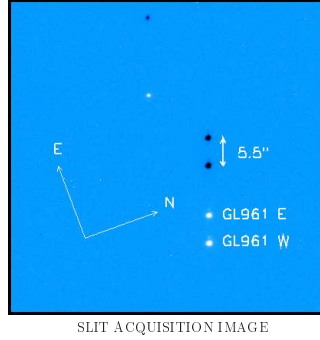


Fig. 2. VLT-ISAAC spectra obtained toward GL 961 E and W. The $3.3 \mu\text{m}$ PAH feature may be affected by poor cancellation of an atmospheric feature.

The L- and M-band windows include the following main features: $3 \mu\text{m}$ (H_2O ice), $3.47 \mu\text{m}$ (unidentified), $3.54 \mu\text{m}$ (CH_3OH ice), $4.08 \mu\text{m}$ (HDO ice), $4.62 \mu\text{m}$ (OCN^- ice) and $4.67 \mu\text{m}$ (CO gas and ice). The strong solid H_2O and CO bands are detected in most objects, although with varying amounts. Typical LR and MR spectra are shown in Figures 1–4. Because of the strong atmospheric features at the relatively low altitude of Paranal, the quality of the data around the solid OCS feature at $4.9 \mu\text{m}$ is low.

Several of these features are excellent indicators of the thermal history and energetic processing of the ices. Specifically, we can use the following diagnostics of the physical conditions and evolution of our sources: (i) the solid H_2O profile, where the peak position gives an indication of the ice temperature; (ii) the solid CO profile, where the shape indicates the ‘apolar’ vs. ‘polar’ ice fraction; (iii) the $\text{CO}/\text{H}_2\text{O}$ abundance ratio, with the more volatile CO molecule having a lower ice abundance at high temperature; (iv) the gas/solid CO ratio and the gas-phase CO excitation temperature; and (v) the presence of the OCN^- feature, which is thought to be a tracer of energetic processing (ultraviolet irradiation or

particle bombardment) (Schutte & Greenberg 1997, Demyk et al. 1998). In the following, a few specific initial results are presented.

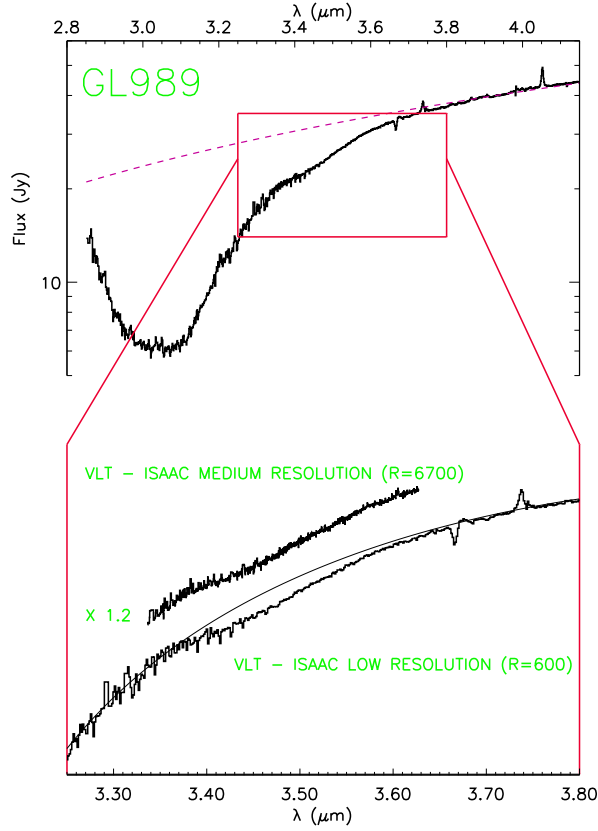


Fig. 3. VLT-ISAAC LR and MR L-band spectra obtained toward GL 989. Note the lack of substructure in the 3.47 μm feature (Dartois et al., in prep).

3.1 The 3.47 μm feature

The 3.47 μm feature, previously detected by Allamandola et al. (1992) and Brooke et al. (2000), is observed in several of our sources at high S/N . The comparison of the LR and MR spectra for the bright source GL 989 shows good agreement in the shape of the spectra (Figure 3). Also, no substructure is apparent at the higher spectral resolution. The data are currently being compared with different laboratory ice mixtures, in particular mixtures involving H_2O and NH_3 (Dartois et al., in prep.). Theoretical models predict that NH_3 is an important component of interstellar ices formed by hydrogenation of atomic N, but unfortunately the strongest NH_3 bands are blended with H_2O at 3 μm and with

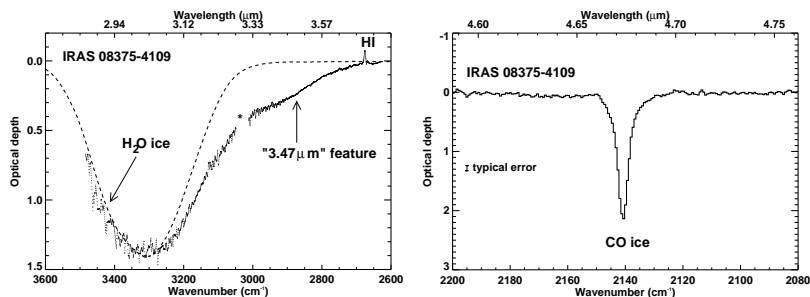


Fig. 4. VLT-ISAAC LR L- and M-band spectra obtained toward IRAS 08375 –4109 in the Vela molecular cloud (Thi et al., in prep).

the silicate band at $9.6 \mu\text{m}$. A tentative detection of the $9.6 \mu\text{m}$ feature toward one (northern) object has recently been claimed by Lacy et al. (1998), suggesting high NH_3 abundances up to 10% with respect to H_2O ice. The analysis of our VLT data indicates lower NH_3 abundances.

3.2 Vela sources

LR L- and M-band spectra of 5 intermediate mass YSOs ($\sim 300 - 700 L_\odot$) have been obtained in the previously unexplored Vela molecular cloud (Thi et al., in prep.). The objects were selected from the list of ‘class I’ objects of Liseau et al. (1992), based on their spectral energy distribution. For two sources, an additional object was found in the field within a few ″, providing ‘off source’ information on the ices. MR spectra in the region of the solid and gas-phase CO band have been taken as well for a few cases.

H_2O ice has been detected in all objects, and CO ice in half of the objects. The H_2O profile indicates that the bulk of the ice is very cold. The solid CO band toward IRAS 08375 –4109 is one of the strongest and sharpest CO bands observed in any source: at LR, the feature is unresolved and its true depth can only be obtained from the MR spectrum (Figure 4). In spite of the low overall temperature, clear differences in the solid CO/ H_2O ice abundances are observed, which correlate with the bolometric temperature of the source.

Several of the sources show the $3.47 \mu\text{m}$ feature, but the presence of the $3.54 \mu\text{m}$ band is less clear. This limits the solid CH_3OH abundance to a few % of that of H_2O ice, significantly less than the 50% found toward some massive protostars (Dartois et al. 1999). The OCN^- feature is detected in at least one high temperature source.

3.3 The circumstellar disk around L1489

L1489 has been shown by Hogerheijde (2001, this volume) to be a transitional object between the class I and II phases. It is surrounded by a large 2000 AU

radius rotating circumstellar disk, which must be on the verge of shrinking to the ~ 100 AU size disks seen around T Tauri stars. Thus, it provides an excellent opportunity to probe the chemical composition of the gas and dust just when it is being incorporated into the disk. The VLT spectra of L1489 show strong absorption by H_2O and CO ices, but no evidence for CH_3OH and OCN^- features, providing limits on their abundances of a few % with respect to H_2O ice.

In summary, the initial data show that the VLT-ISAAC is a powerful instrument to obtain high-quality $2.9\text{--}5\ \mu\text{m}$ spectra of low-luminosity embedded YSOs, and that such data can provide an important step forward in our understanding of the physical and chemical evolution of ices in low-mass young stellar objects and their incorporation into new planetary systems.

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